**Euclid**

Mission Overview

The Euclid mission is planned for launch in 2018redbook page 8 2012 version, at an estimated total cost of 800 million Euros, and will work in the optical and near infrared. http://www.bbc.co.uk/news/science-environment-18503703. Its primary goal is to conduct a wide survey – some 15000 degrees of sky is planned to be covered – in order to map the geometry of the dark universe. It will look at galaxies and galaxy clusters to work out their redshifts and shapes, back to z= 2. There is also to be a deep survey which is expected to cover around 40 degrees of sky to a depth 2 magnitudes deeper than the wide survey. This deep survey will allow astronomers to see even further back in time, up to redshifts of 8 and potentially even higher. The primary mission objectives are expected to be completed within 7 years.

One of Euclid’s main scientific objectives with the deep survey is to study high redshift galaxies at z =6+ over a very wide survey area. This will give astronomers the opportunity to spectroscopically confirm hundreds of galaxies for use in the study of the EoR. It will help constrain the bright end of the luminosity function at high z. Euclid will ‘*detect hundreds of z=7 galaxies brighter than an apparent magnitude J=26 and tens at z greater than 8.*’ Euclid Definition Study Report 2012 Published 29 September 2011 R. Laureijs et al

Another objective is to conduct a weak gravitational lensing survey (as discussed in section ?? Grav lensing by Catherine ) which will be looking for disturbances and distortions in the light from galaxies as a result of large amounts of mass perturbing the path of the photons emitted from them. ‘*Euclid measures the shapes of 30 resolved galaxies per arcmin2 in one broad visible R+I+Z band (550-920 nm) down to AB mag 24.5*’ Euclid Definition Study Report 2012 Published 29 September 2011 R. Laureijs et al PAGE 9

Capabilities

* Visual Imaging/ Photometry, 550 – 900nm
* Spectroscopy, 1100 – 2000nm
* NIR Imaging/ Photometry, 920 – 2000nm (Y, J,H bands)

Euclid will have two instruments in order to do the above; a wide-band imaging system in the visible (VIS), and an instrument capable of both slitless spectroscopy as well as NIR imaging (NISP). These instruments will be able to operate simultaneously when required.

Method of conducting the deep survey

The Euclid deep survey will be carried out by re-visiting particular areas in the wide survey over an extended period of time. Over this time, the depth will be built up by combining a series of images taken of the same piece of sky (known as stacking). The wide survey will reach the required magnitude depth with one exposure, but to reach 2 magnitudes deeper (26th magnitude AB at 5 sigma) considerably more exposures will be needed. Figure (No. OF VISITS…) shows the number of visits to increase the depth to 26th magnitude. The current estimate is that this will take around 36 exposures, but it has been decided that 40 will be done to guarantee that the depth requirement is met.

NISP performance analysis report Ref : EUCL- MPI-NPS-RP-00079  
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**Figure ?: Number of visits required to meet the Euclid’s goal of a depth of 26th magnitude**

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Uses and Limitations for Euclid’s use in studying high redshift galaxies

The main limitation of using Euclid for this study is its filter range: Since the drop out technique requires two filters above the drop and one below, the redshift of the galaxy must be such that the filters available are able to achieve this. Euclid’s longest wavelength filter is centred at 1.63 microns (J-band filter). If the drop is at too long a wavelength, only one band will observe flux. Without two bands observing flux, no colour-colour diagrams can be plotted, and one measurement of flux and another of no flux is not enough to have confidence that the object is a Lyman break galaxy. Table ? shows the IR filters of Euclid, along with each filter’s central wavelength and it’s bandwidth.

|  |  |  |
| --- | --- | --- |
| filter | Central wavelength | bandwidth |
| Y (920 – 1146) | 1020 | 226 |
| H (1146 – 1372) | 1220 | 226 |
| J (1372 – 2000) | 1630 | 628 |

From this, it can be concluded that the survey will be unhelpful in determining galaxies past a redshift of 10.5 (corresponding to the drop being observed at 1.40microns): The object would be detected in the J band, but not in either H or Y. The end of the H filter and start of the J filter is at 1.37 microns. To get significant flux in the Y band, ideally the galaxy would need to be at a redshift of 10 or lower, such that a noticeable amount of flux was detected in H.

The weak lensing bit that was here has been deleted.

The spectroscopic mode can be used (grism spectroscopy) to determine the redshift of LBGs as outlined in section (john’s spectroscopy section). The advantage of grism spectroscopy is that multiple objects can be studied simultaneously, meaning if there are multiple likely candidates in the field of view, their spectra can all be gathered at once, reducing the overall time taken to study a large number of redshift galaxies. This is particularly useful to Euclid, which has a wide field of view of just over 0.5 square degrees, increasing the chance of multiple candidates.

Key technical data:

|  |  |
| --- | --- |
| Primary mirror | 1.2m |
| FoV | 0.763x0.763 degrees2 |
| Pixel size | 0.3 arcsec X 0.3 arcsec |
| Detector Array | 2kx2k pixels |
| Resolution | 0.3 -> 0.6 arcsecs (in J band) |
| Plate Scale (infrared) | 0.3arcsec per pix |

The data is quoted for the deep survey NIR photometry. Some data is subject to slight change as the planning stages progress.